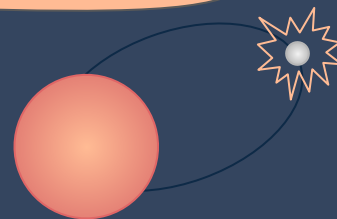
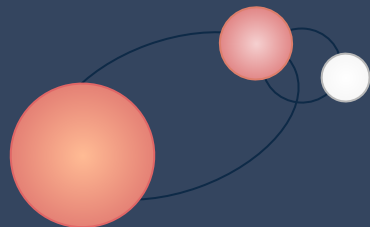
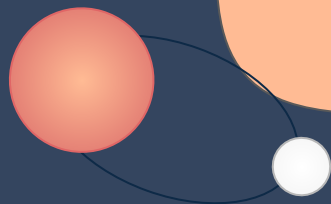


56 Ursae Majoris: a binary evolution puzzle



Ana Escorza

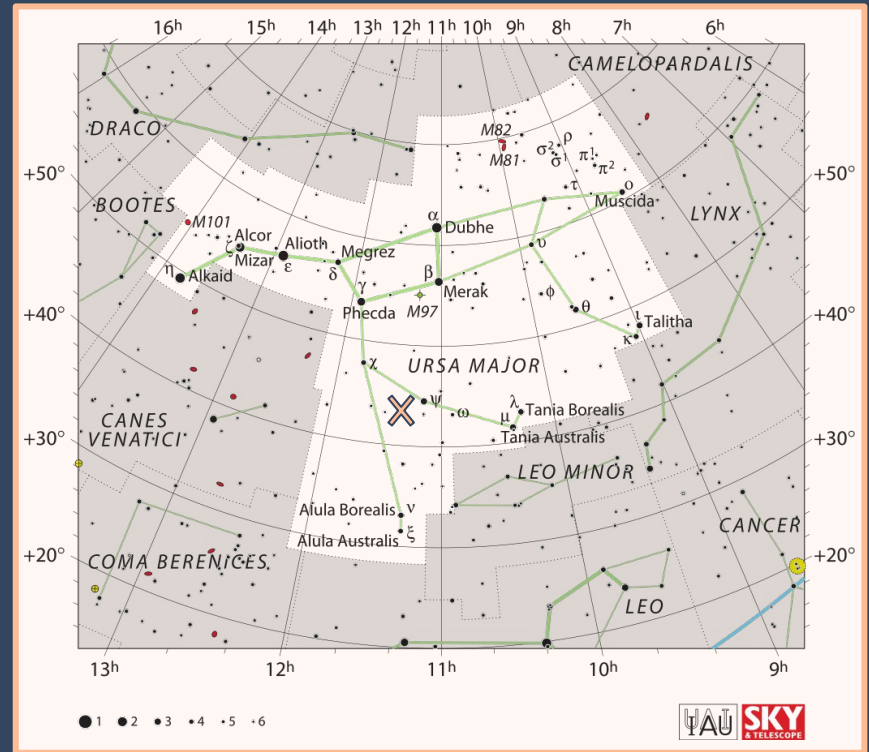
with D. Karinkuzhi, A. Jorissen, S. Van Eck,
J. T. Schmelz, G. L. Verschuur, H. M. J. Boffin,
R. J. De Rosa, and H. Van Winckel



"la Caixa"

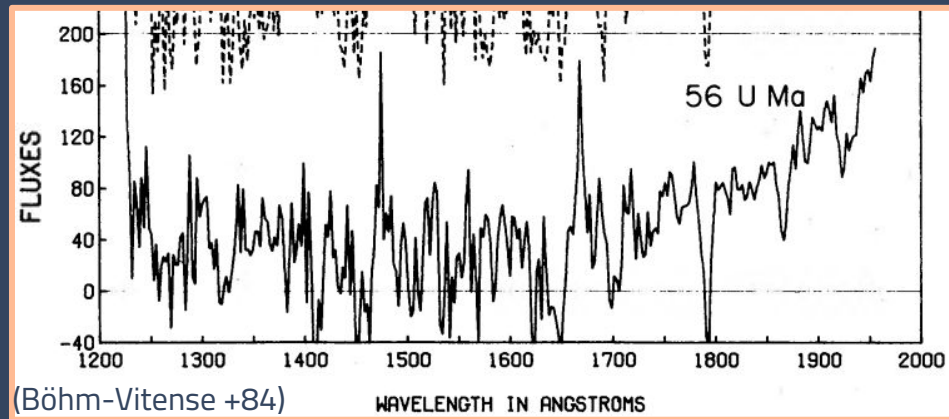
56 Ursae Majoris in the literature

- G8III + faint lower-mass companion (Keenan & McNeil 1989)
- $V = 5.03$ mag (Ducati 2002)
- Strong detection in the Galex far-UV band (Bianchi et al. 2011, 2017)



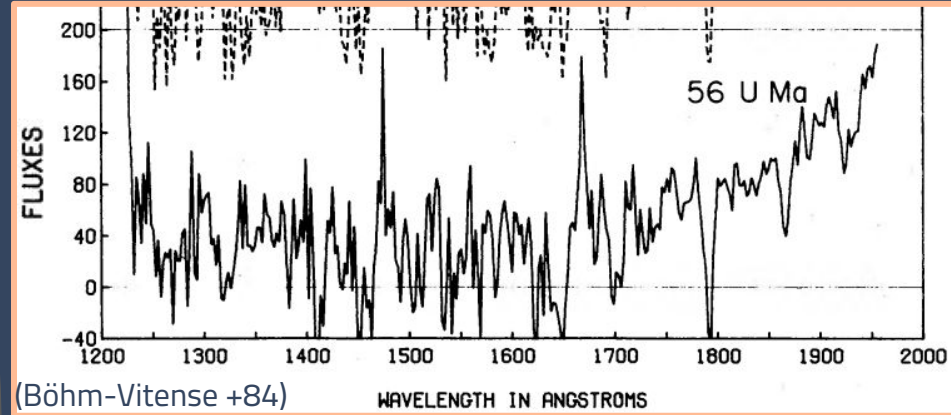
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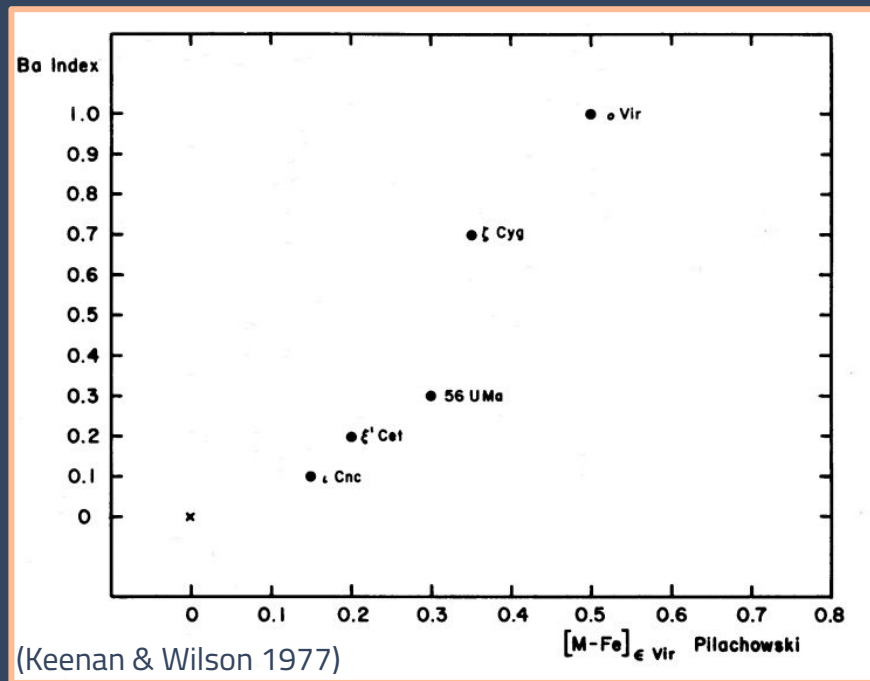
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56 Ursae Majoris as a Barium star

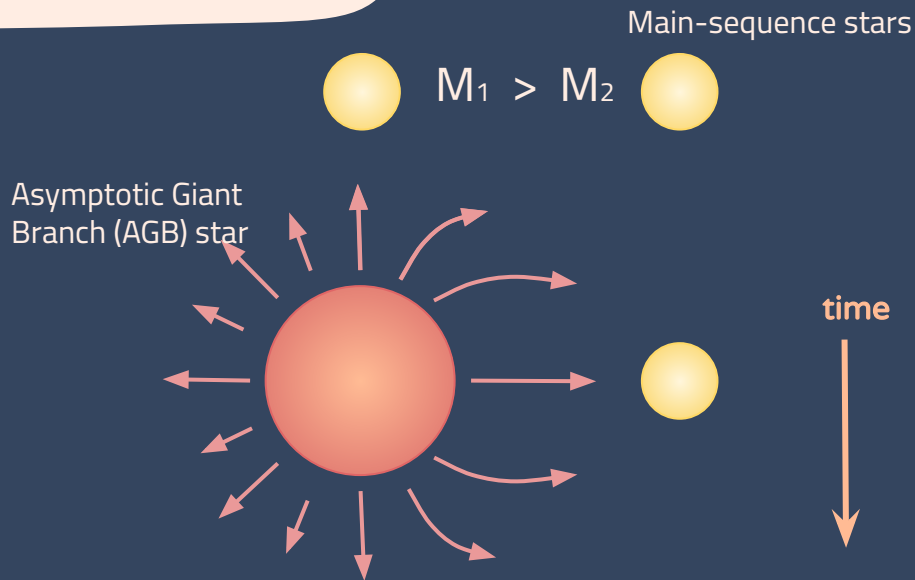
Main-sequence stars



- G8III + faint lower-mass companion (Keenan & McNeil 1989)
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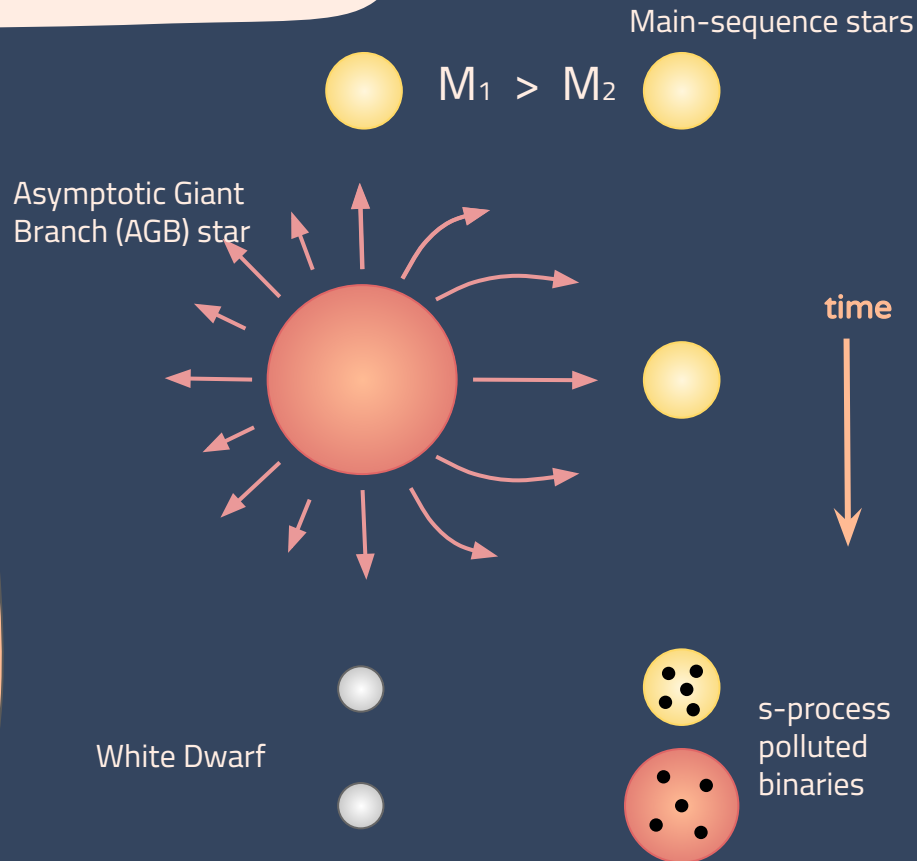
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56 Ursae Majoris as a Barium star

- G8III + faint lower-mass companion (Keenan & McNeil 1989)

- $V = 5.03$ mag

- Strongly enriched in barium (Bianchi et al. 1991)

- Two white dwarf companions detected (Böhm-Vitense et al. 1997)

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- The giant was classified as a mild barium star (Pilachowski 1975; Keenan & Wilson 1977)

If the giant is a Ba star,
the companion MUST be a white dwarf

Asymptotic Giant Branch (AGB) star



White Dwarf



s-process
polluted
binaries



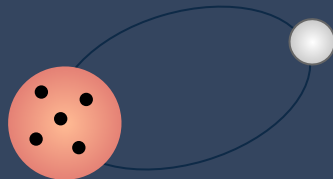
time



The masses of the WD companions of Ba stars

Combining RV data with Hipparcos and Gaia astrometry and the Hip-Gaia catalogue of accelerations, we could constrain orbital inclinations and secondary masses for 60 WD companions of Ba stars

(Escorza & De Rosa 2023)



Hipparcos



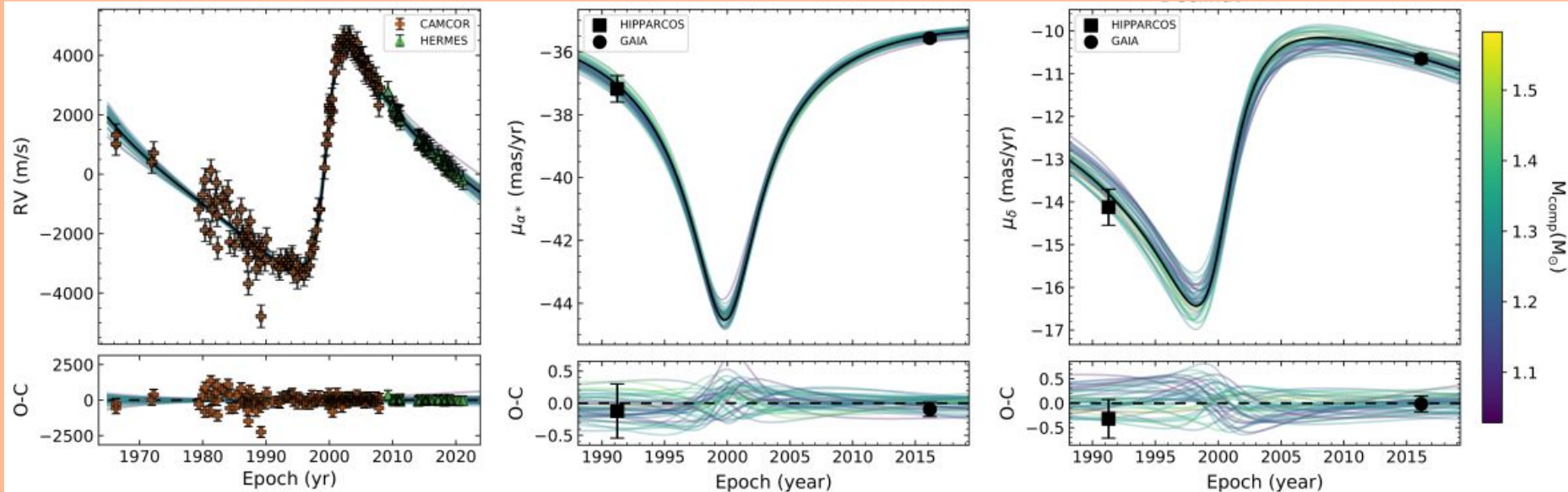
Gaia

The masses of the WD companions of Ba stars

RV data + Hipparcos-Gaia acceleration

work done using ORVARA (Brandt et al. 2021)
and the HGCA (Brandt 2018, 2021)

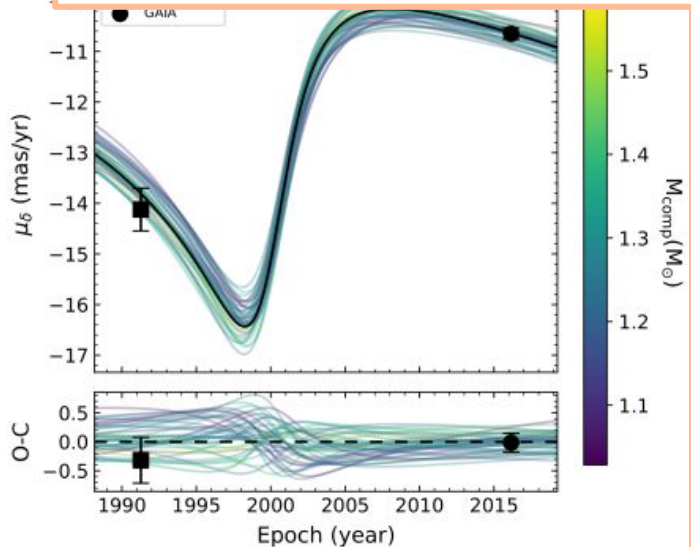
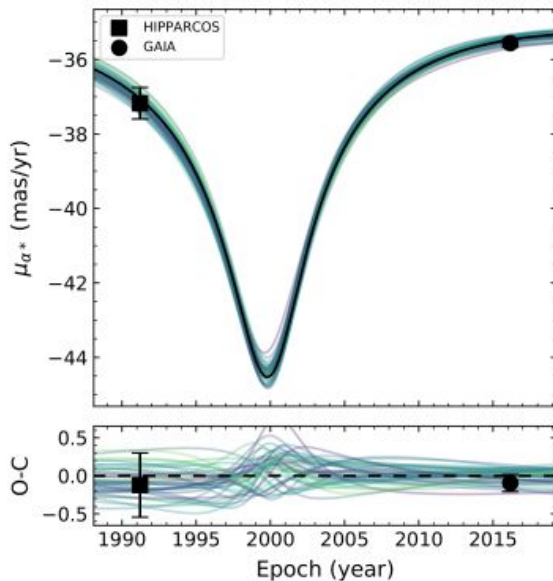
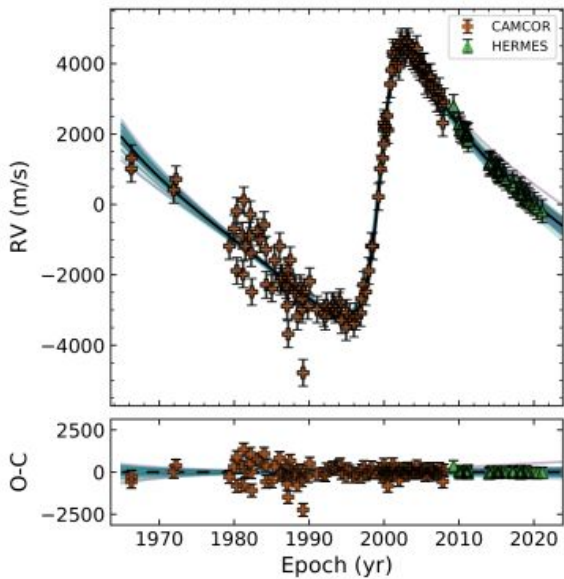
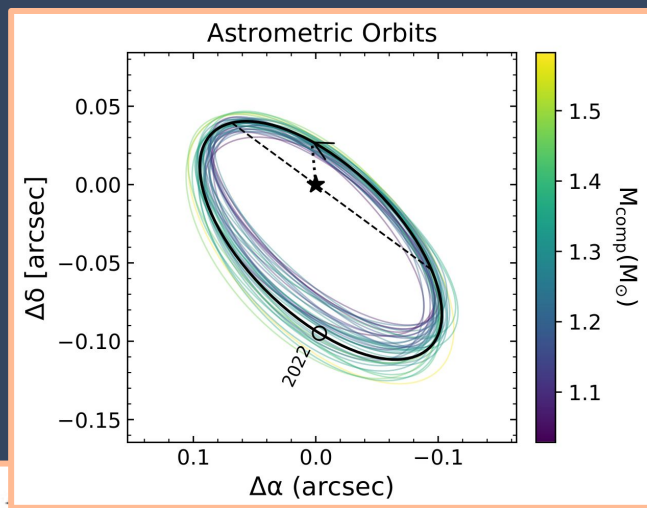
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The masses of the WD companions of Ba stars

RV data + Hipparcos-Gaia acceleration
+ Hipparcos epoch astrometry

(Escorza & De Rosa 2023)

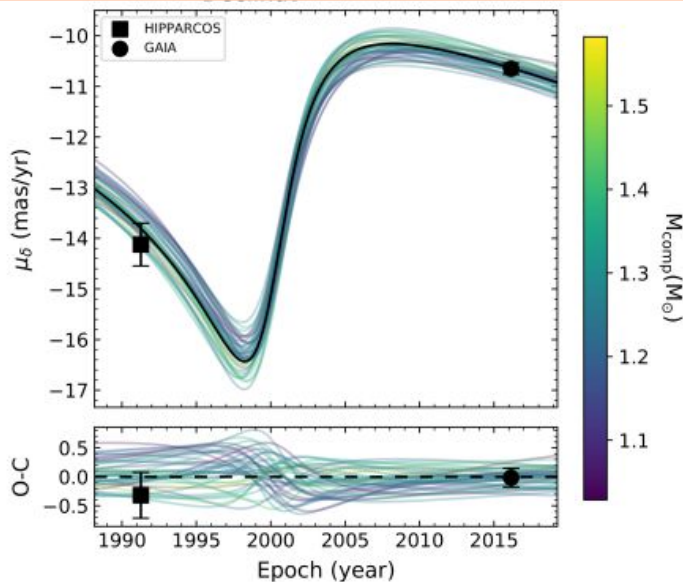
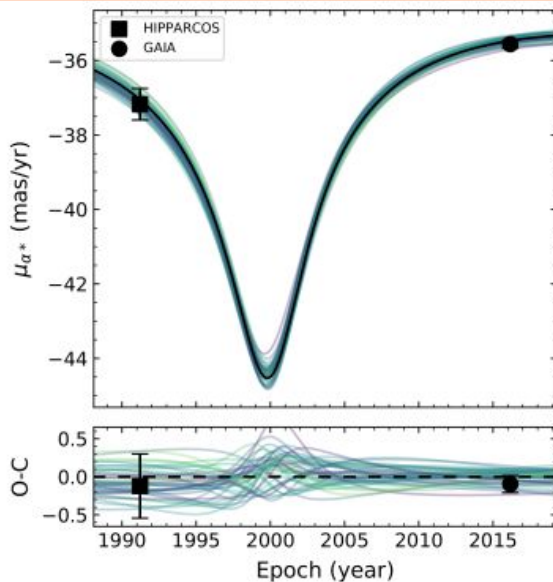
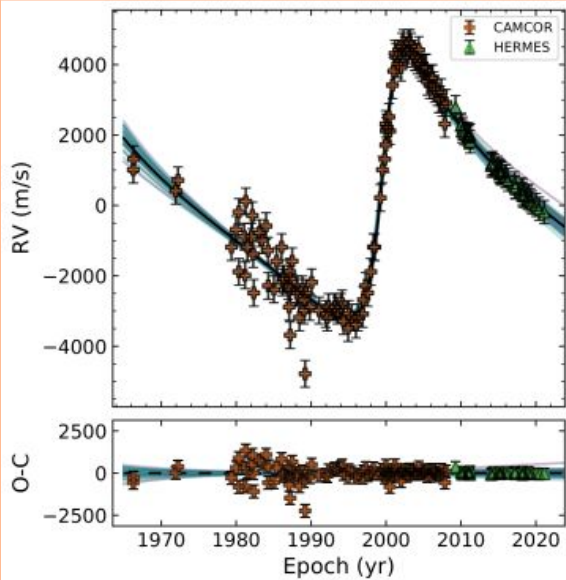


The masses of the WD companions of Ba stars

RV data + Hipparcos-Gaia acceleration

+ Hipparcos epoch astrometry \Rightarrow 56 UMa has a $1.31 \pm 0.12 M_{\odot}$ companion

(Escorza & De Rosa 2023)

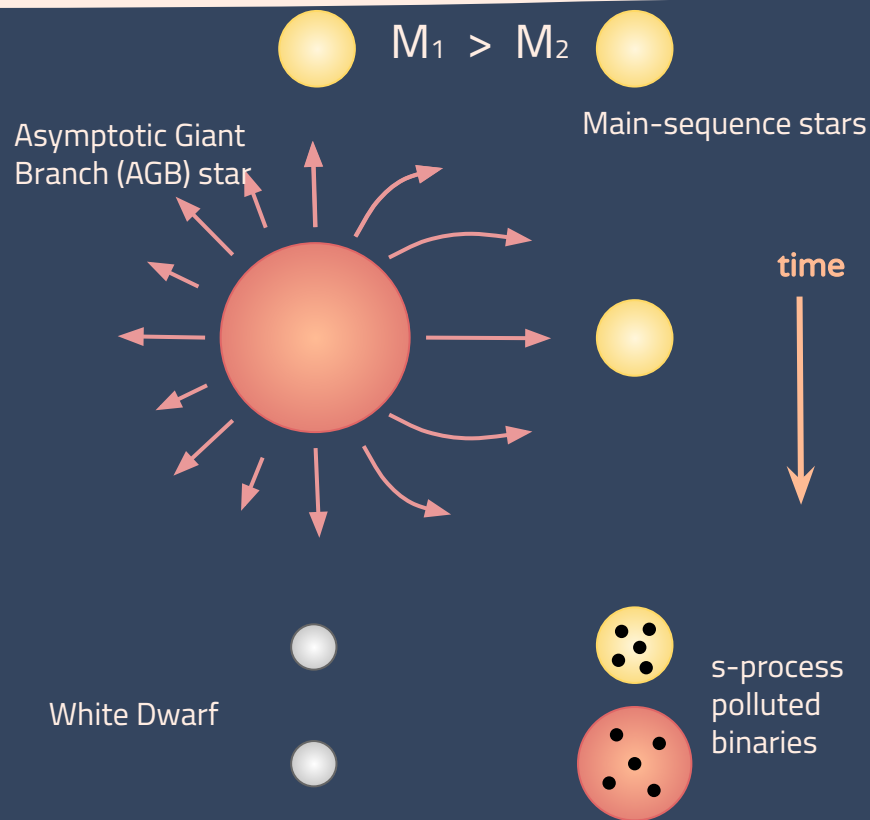


The mass of the faint component in 56 UMa

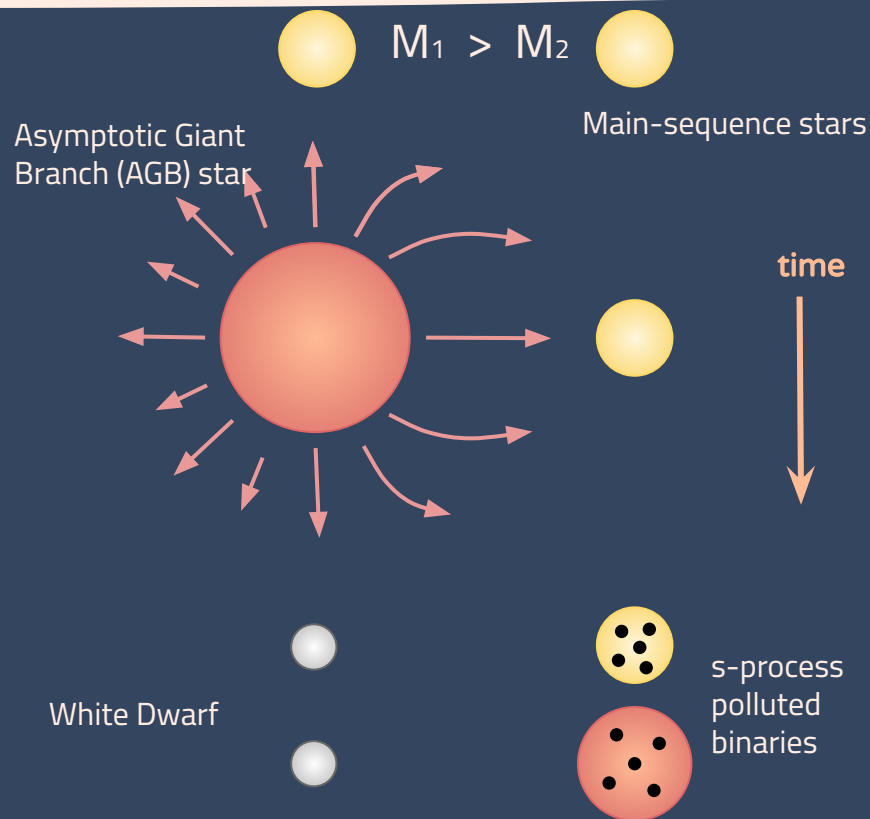
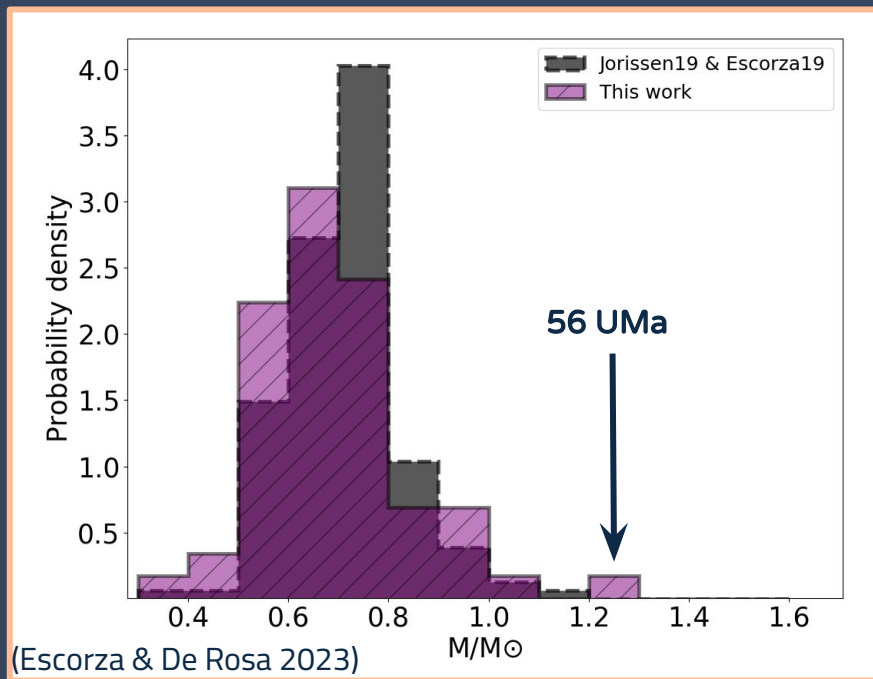
56 UMa has a $1.31 \pm 0.12 M_{\odot}$ companion

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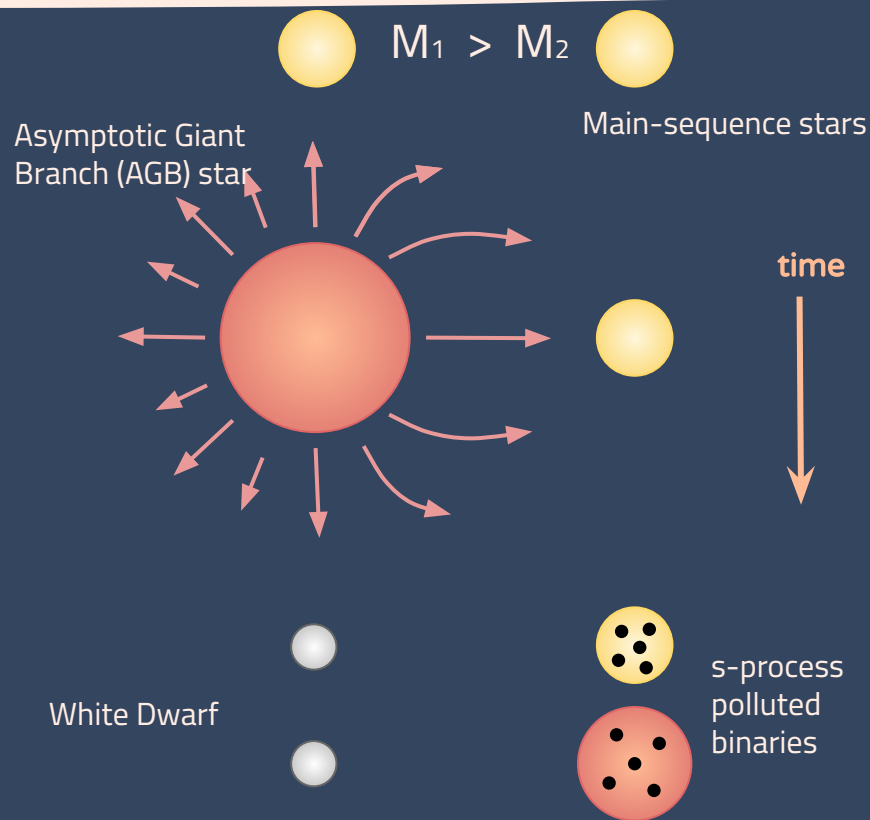
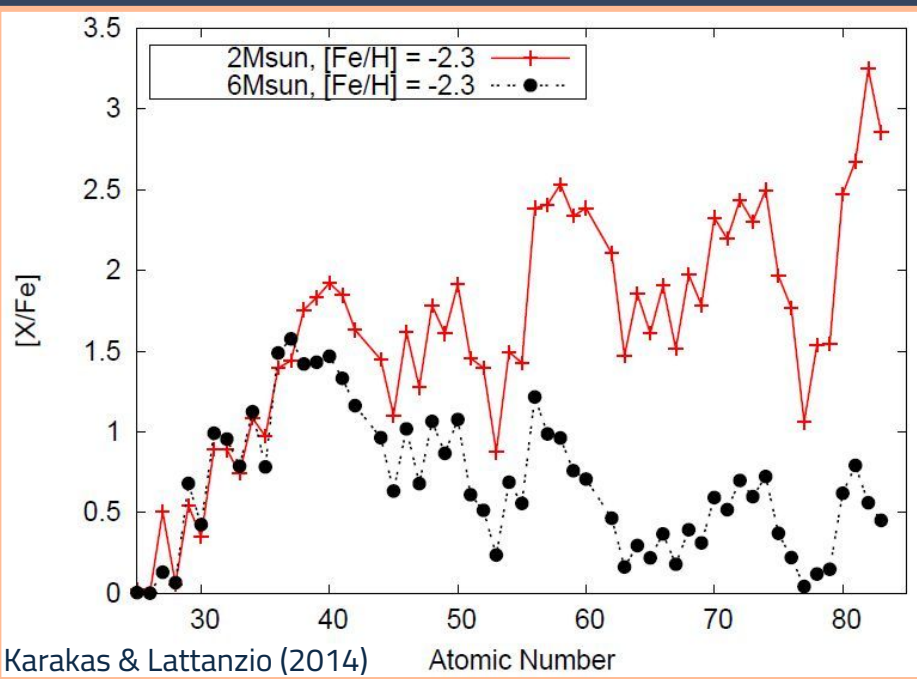


The mass of the faint component in 56 UMa



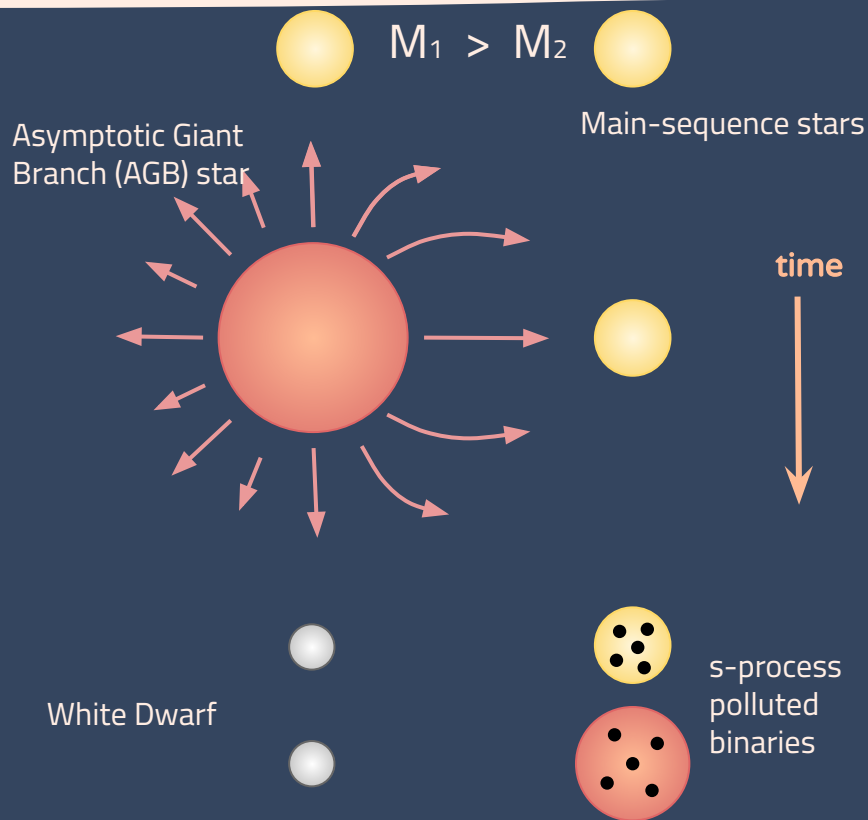
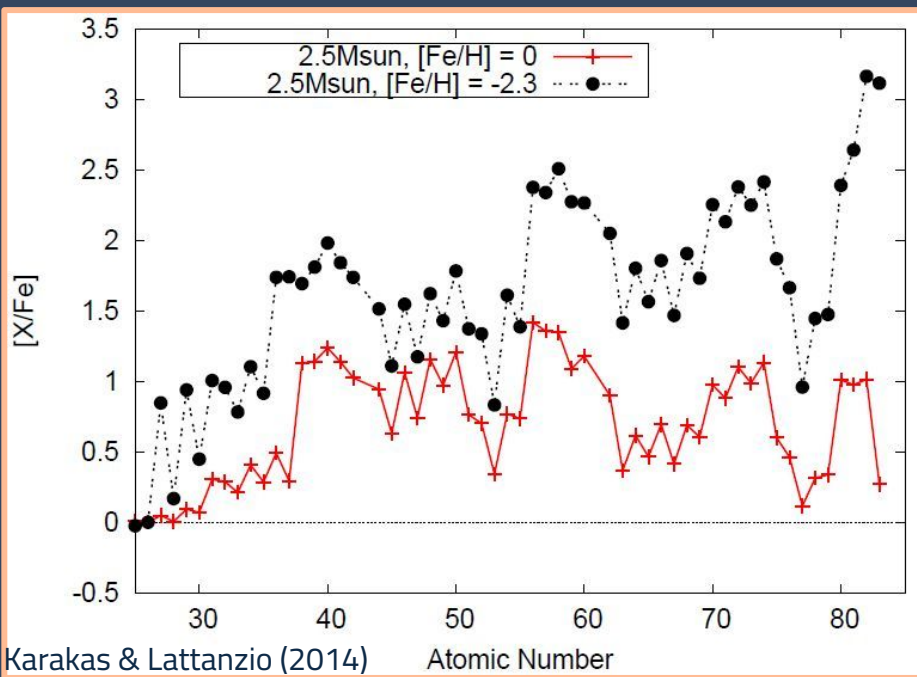
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Could the AGB progenitor of such a WD have made a Ba star?

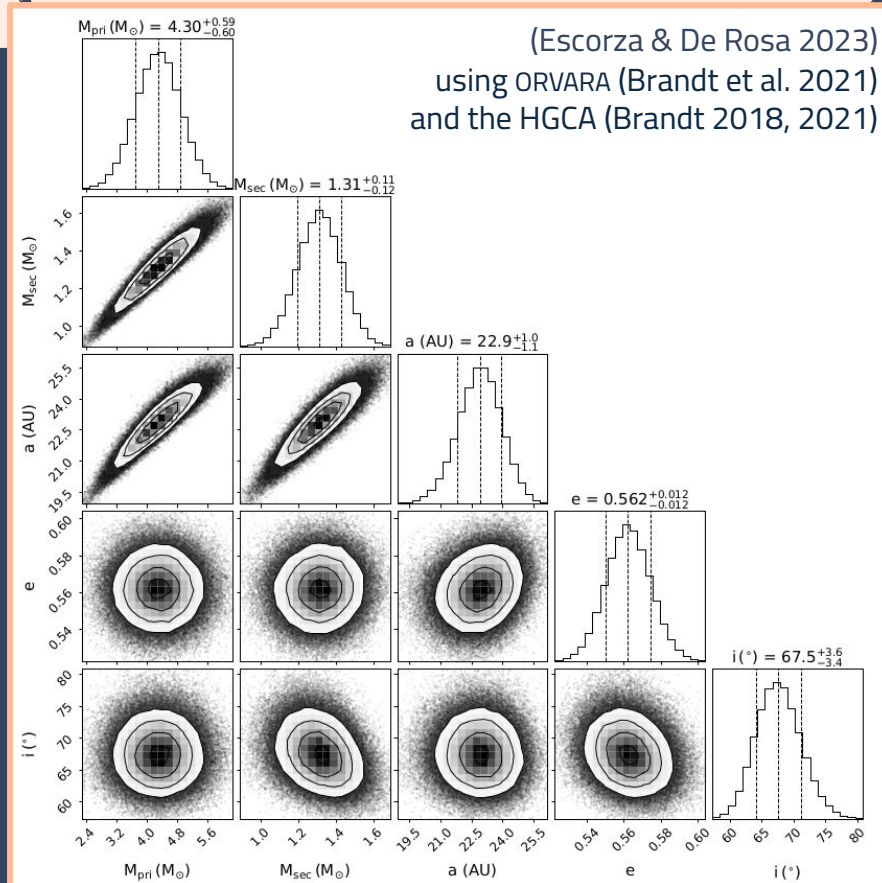


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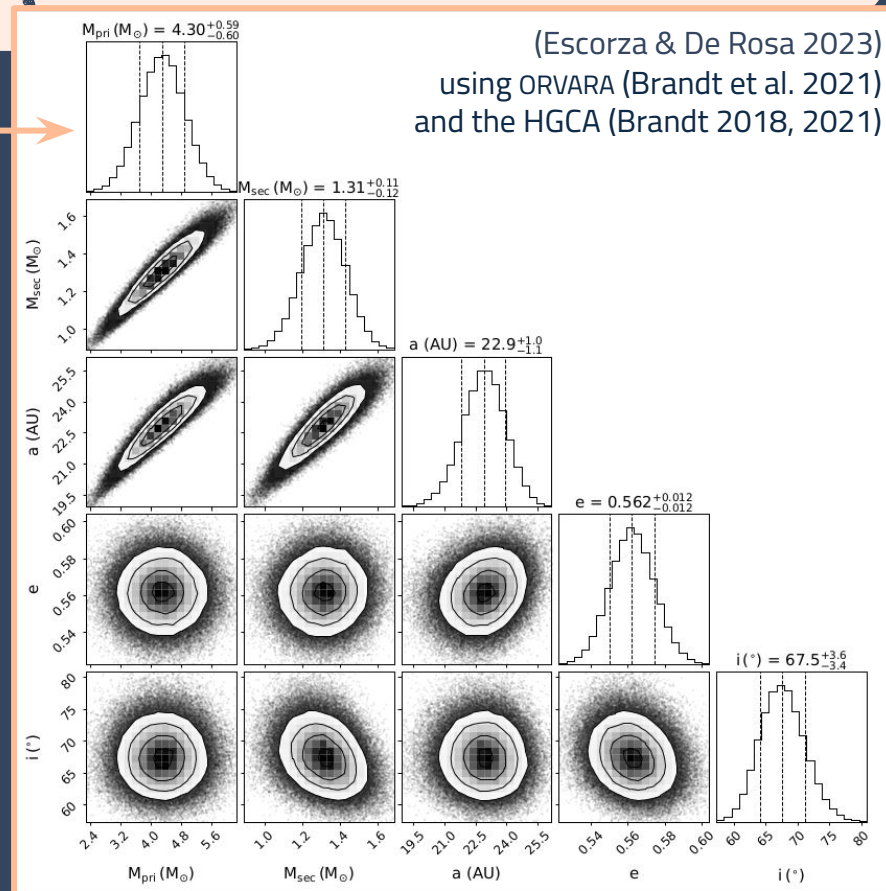


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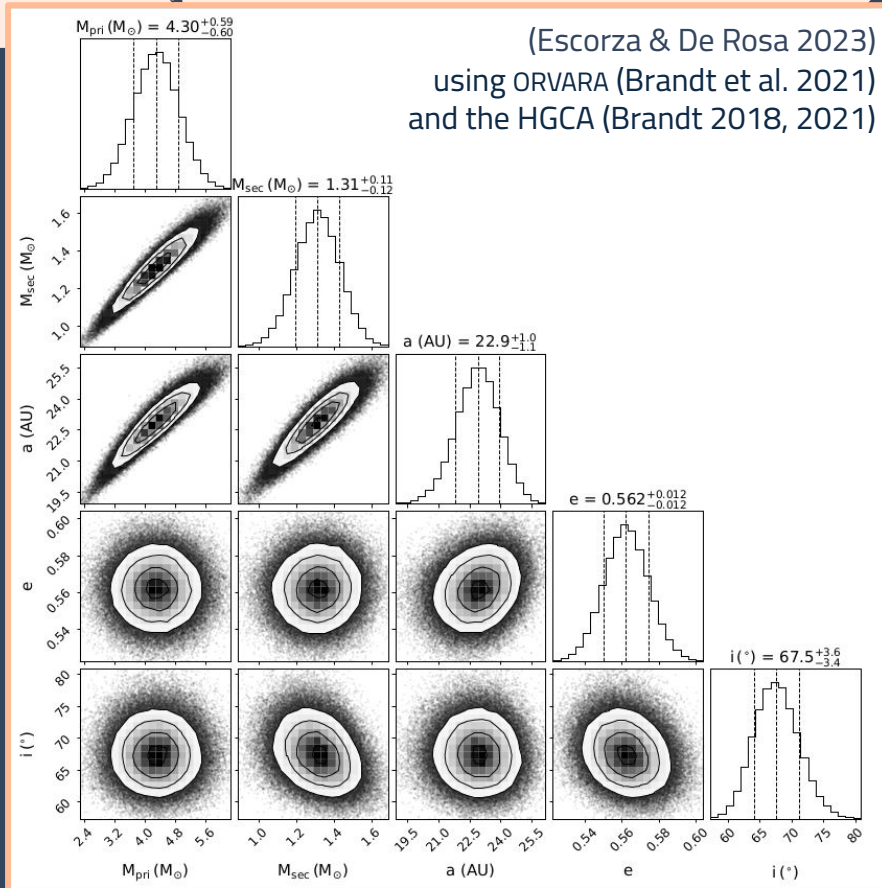
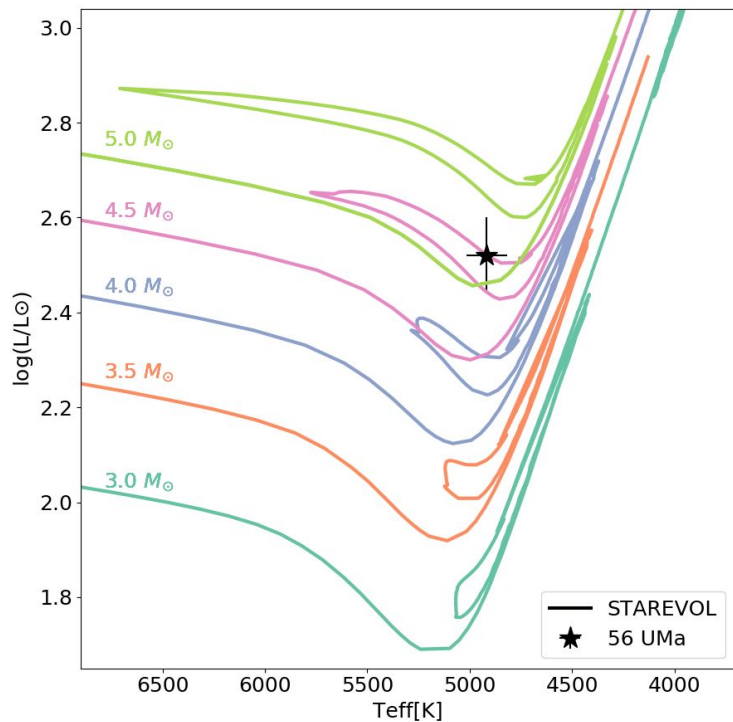
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Is the primary mass correct? →



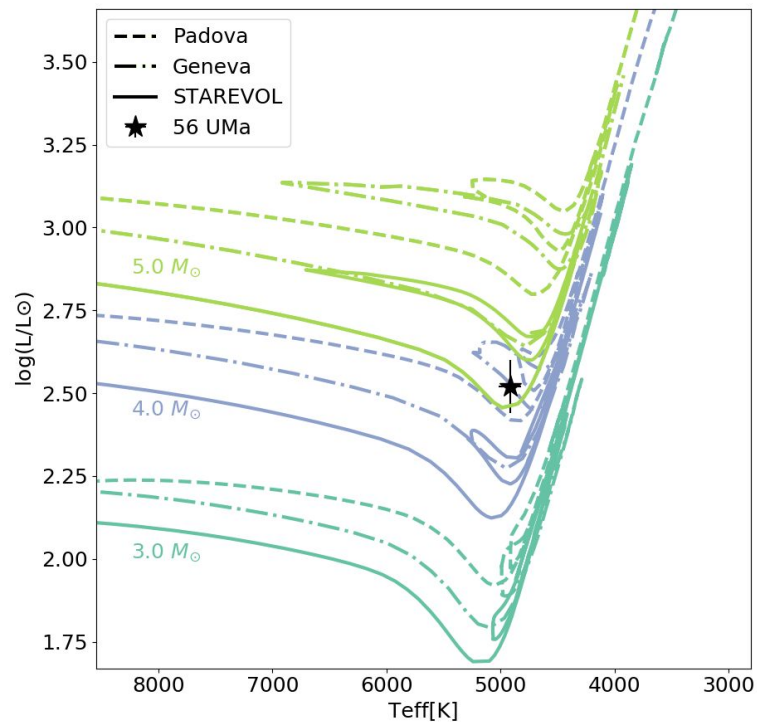
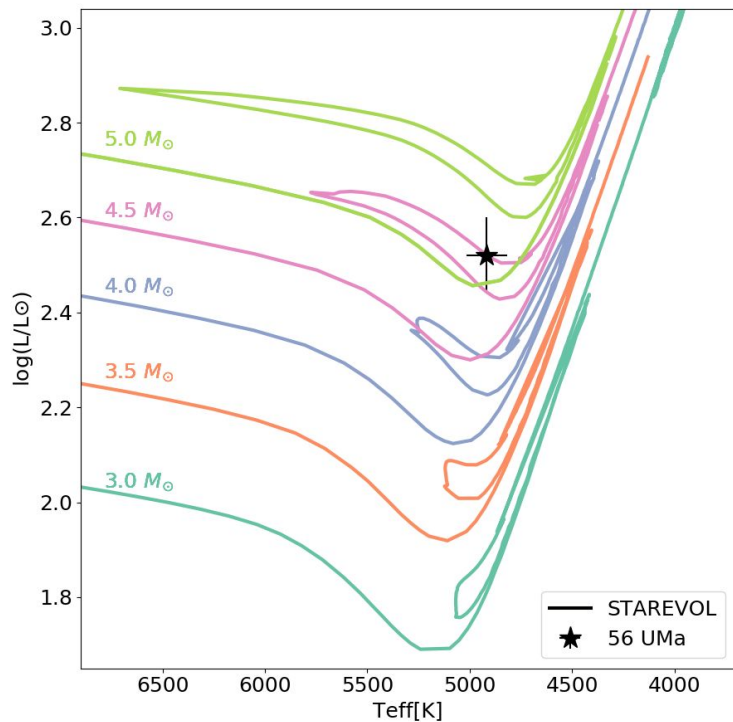
The mass of the red giant component in 56 Uma

Jorissen et al. (2019); Escorza et al. (2023)

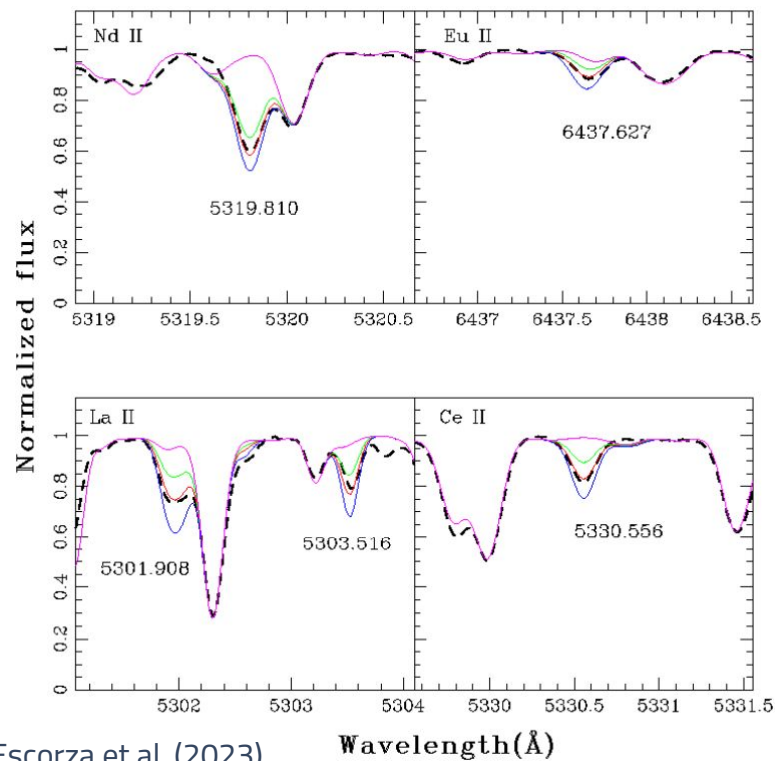


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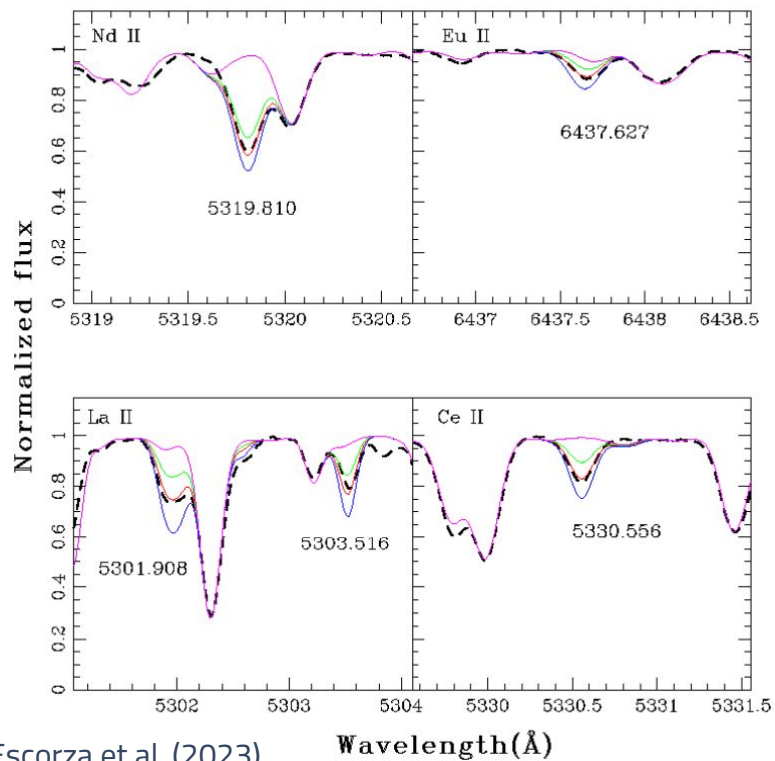
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The heavy metal abundances of the red giant

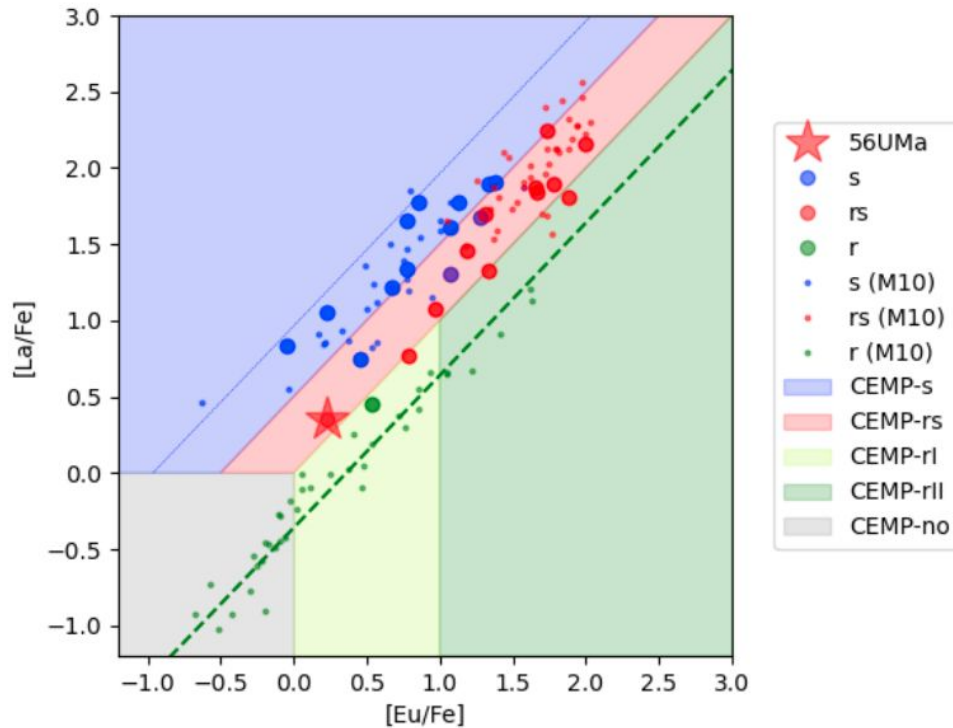


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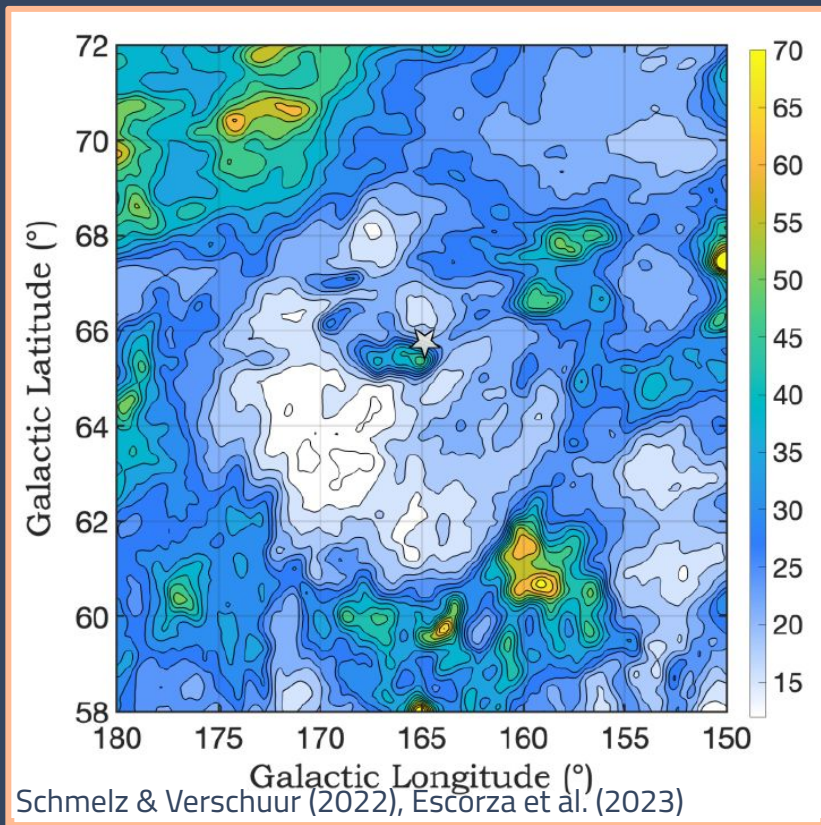
Escorza et al. (2023)

Wavelength(Å)



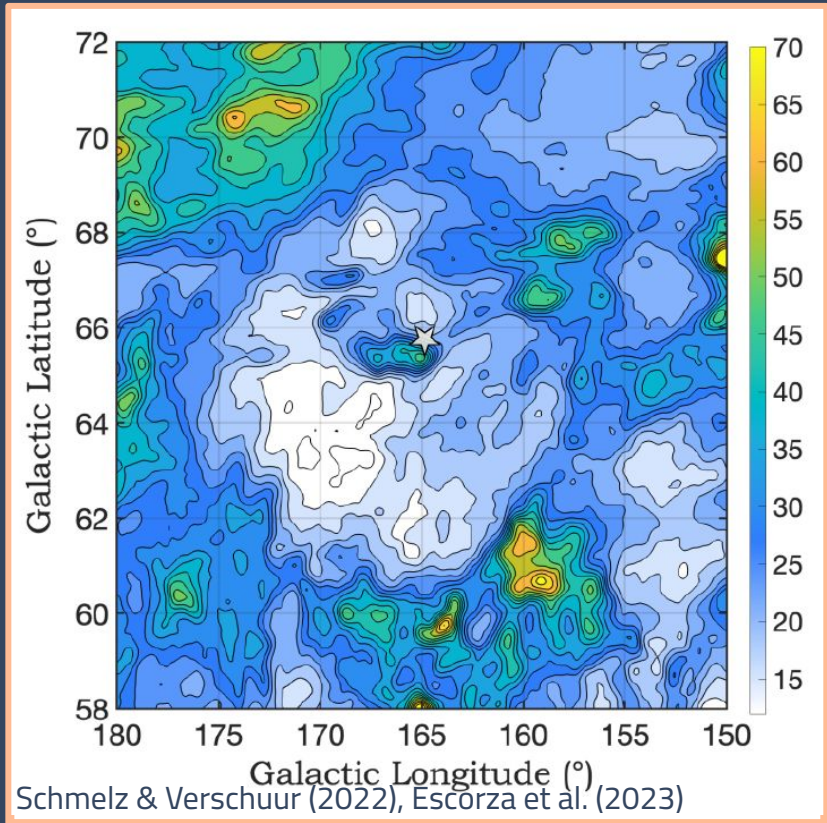
Escorza et al. (2023); comparison CEMP stars from Karinkuzhi et al. (2021)

The neutral gas cavity around 56 UMa



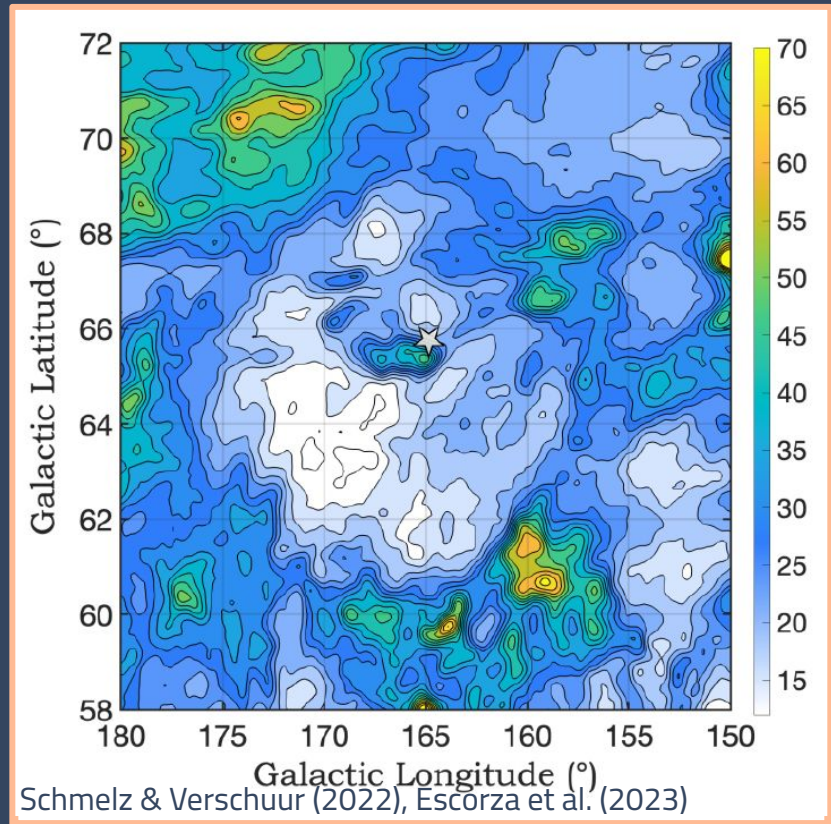
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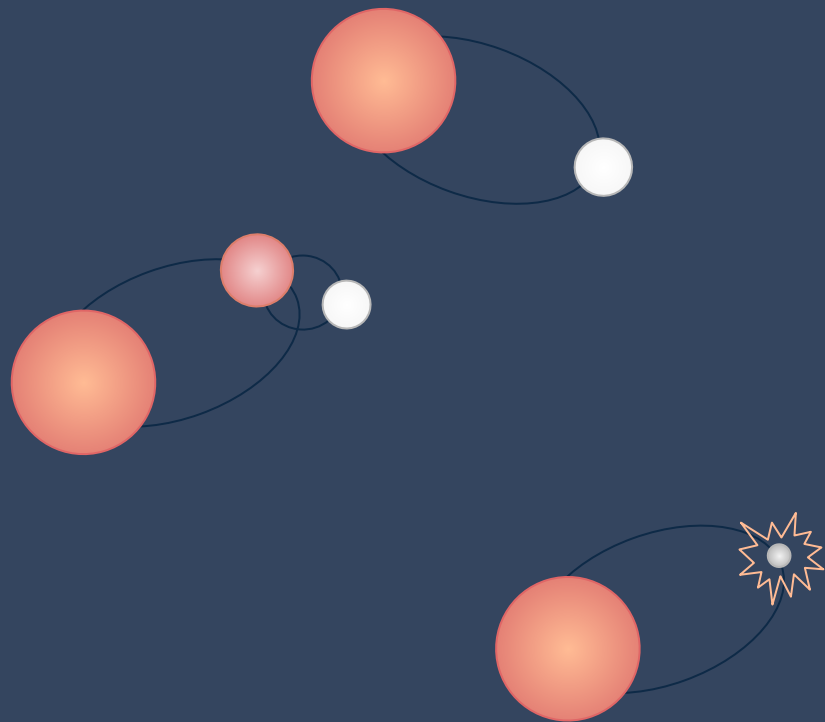
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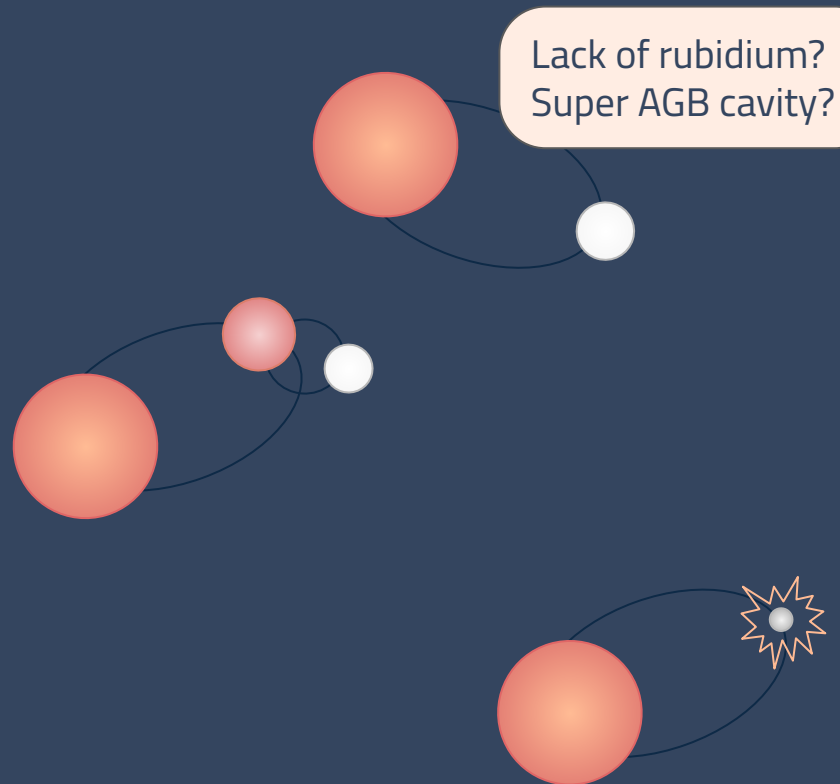
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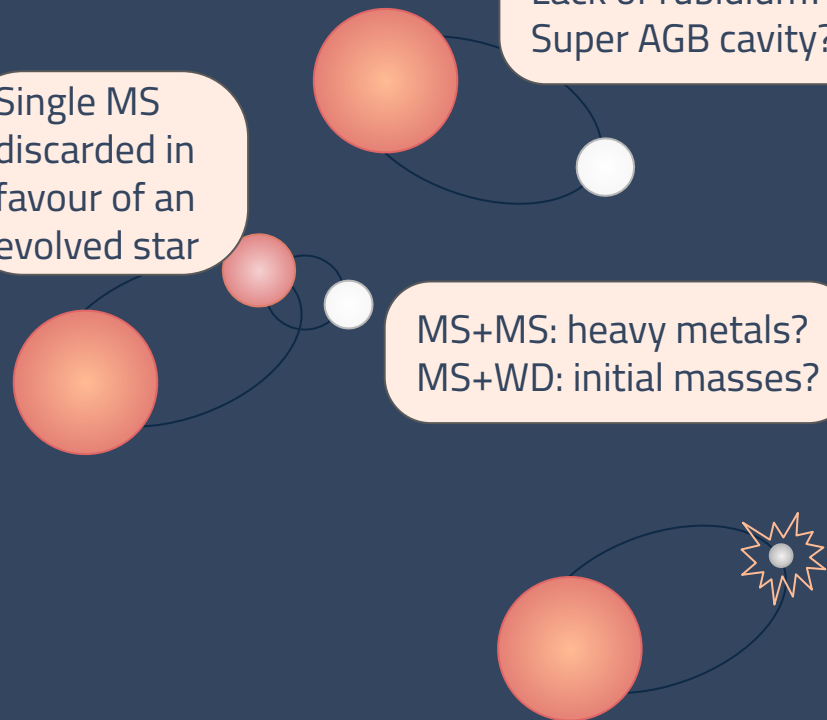
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Single MS
discarded in
favour of an
evolved star

Lack of rubidium?
Super AGB cavity?

MS+MS: heavy metals?
MS+WD: initial masses?



The nature of the faint component in 56 UMa

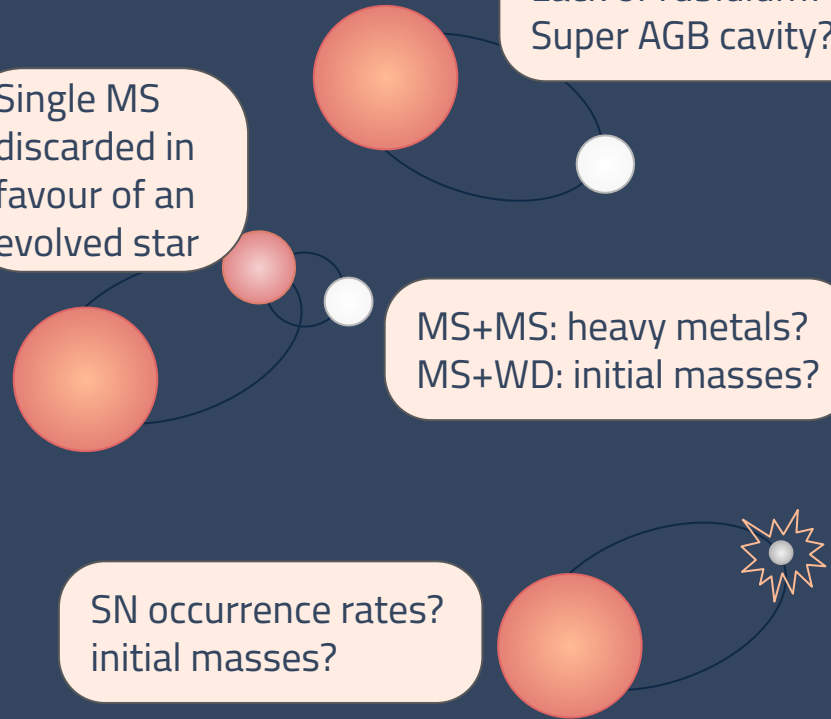
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SN occurrence rates?
initial masses?



Thanks!

Table 1. Stellar and orbital parameters of the 56 UMa system.

Parameter	Median $\pm 1 \sigma$
Temperature, T_{eff} [K]	4917 ± 34
Surface gravity, $\log g$	2.3 ± 0.6
Metallicity, [Fe/H]	-0.05
Microturbulence, ξ [km s^{-1}]	1.56
Primary mass, M_1 [M_{\odot}]	4.3 ± 0.2
Period, P [days]	$16\,911^{+438}_{-401}$
Eccentricity, e	$0.562^{+0.012}_{-0.012}$
Semi-major axis, a [AU]	$22.9^{+1.0}_{-1.1}$
Argument of periastron, ω_1 [$^{\circ}$]	$286^{+2.3}_{-2.3}$
Time of periastron, T_0 [HJD]	$2\,468\,401^{+432}_{-385}$
Parallax, ϖ [mas]	$5.86^{+0.03}_{-0.04}$
Ascending node, Ω [$^{\circ}$]	60^{+3}_{-3}
Inclination, i [$^{\circ}$]	$68^{+3.6}_{-3.4}$
Secondary mass, M_2 [M_{\odot}]	$1.31^{+0.11}_{-0.12}$
Center-of-mass velocity [km s^{-1}]	0.13 ± 0.01
Center-of-mass μ_{α^*} [mas yr^{-1}]	-37.32 ± 0.01
Center-of-mass μ_{δ} [mas yr^{-1}]	-12.18 ± 0.01

Table D.1. Individual abundances of the giant component in 56 UMa

	Z	$\log \epsilon_{\odot}^{\text{a}}$	$\log \epsilon$	σ_s (N)	[X/Fe] $\pm \sigma_{[\text{X}/\text{Fe}]}$
C ^b	6	8.43	8.20	0.06(4)	-0.18 ± 0.15
¹² C/ ¹³ C			19		
N ^c	7	7.83	8.40	0.09(30)	0.62 ± 0.21
O ^d	8	8.69	8.70	0.00(2)	0.06 ± 0.22
Na I	11	6.24	6.40	0.10(4)	0.21 ± 0.40
Mg I	12	7.60	7.50:	0.10(2)	-0.05 ± 0.16
Fe I	26	7.50	7.45	0.10(65)	–
Rb I	37	2.52	2.50:	0.00(2)	0.03 ± 0.10
Sr I	38	2.87	3.30:	0.10(1)	0.48 ± 0.35
Sr I _{NLTE}	38	2.87	3.49:	0.10(1)	0.67 ± 0.35
Y II	39	2.21	2.40	0.05(7)	0.24 ± 0.17
Zr I	40	2.58	2.43	0.13(3)	-0.10 ± 0.30
Zr II	40	2.58	2.65	0.06(2)	0.12 ± 0.30
Ba II	56	2.18	2.72:	0.09(2)	0.59 ± 0.11
La II	57	1.10	1.40	0.12(8)	0.35 ± 0.22
Ce II	58	1.58	1.70	0.13(8)	0.17 ± 0.24
Pr II	59	0.72	0.75	0.05(3)	0.08 ± 0.25
Nd II	60	1.42	1.72	0.13(13)	0.35 ± 0.22
Sm II	62	0.96	1.08	0.13(4)	0.17 ± 0.32
Eu II	63	0.52	0.70	0.00(2)	0.23 ± 0.32
Gd II	64	1.10	1.40	0.10 (1)	0.35 ± 0.21
Dy II	66	1.10	1.30	0.10 (1)	0.25 ± 0.18
Hf II	72	0.85	1.10:	0.00(2)	0.30 ± 0.32
Os II	76	1.40	1.60:	0.10 (1)	0.25 ± 0.20